

Origin of the ankle in the UHECR spectrum and extragalactic protons below it

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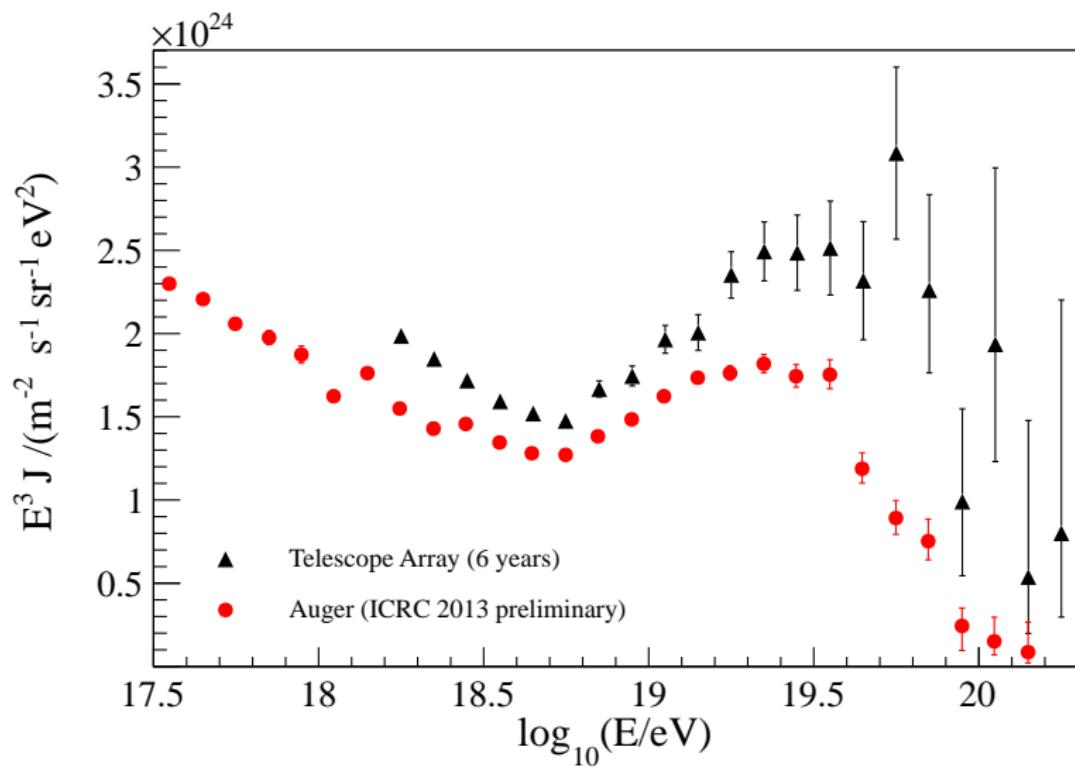
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IceCube Particle Astrophysics Symposium
Cosmic Neutrinos: What's Next?
May 5, 2015

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High energy end of CR spectrum

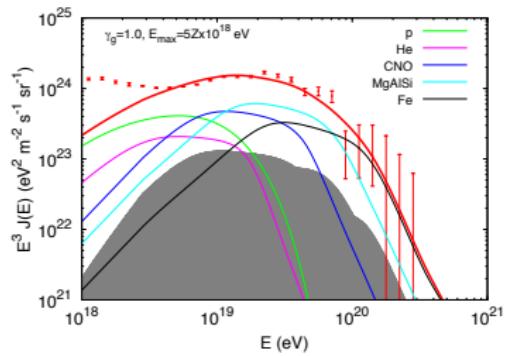


UHECR conundrum

- ➊ Ankle shape readily produced by superposition of two power-laws
Natural candidate: transition between GCRs and EGCRs
 - Original models ↗ transition from Galactic ^{56}Fe to EG protons
(Bird+, 1993)
 - Recent models ↗ transition from G ^{56}Fe to EG heavies
(Allard-Olinto-Parizot, 2007)
- ➋ Ankle feature also naturally arises as dip in spectrum
from e^+e^- energy loss of EG protons propagating in CMB
(Berezinsky-Gazizov-Grigorieva, 2002)
- ➌ Auger data ↗ light but EG component near and below ankle
+ intermediate composition above
(Auger Collaboration, 2014)
- ➍ Recent models ↗ fit Auger spectrum and composition
at price of adding an *ad hoc* light EG component below ankle
with a steep injection spectrum $\propto E^{-2.7}$

(Gaisser-Stanev-Tilav, 2013; Aloisio-Berezinsky-Blasi, 2014)

Example of 4 (Aloisio-Berezinsky-Blasi, 2014)



p-He EG

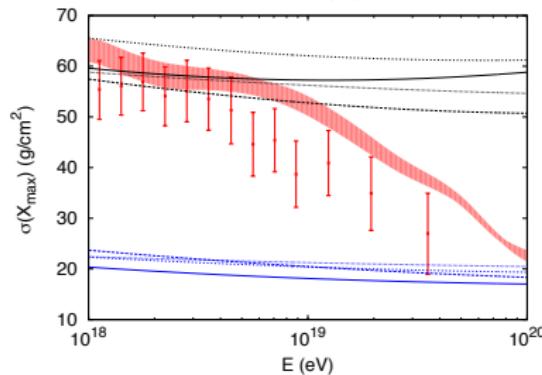
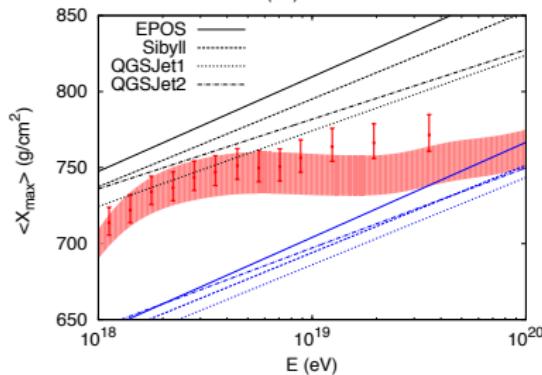
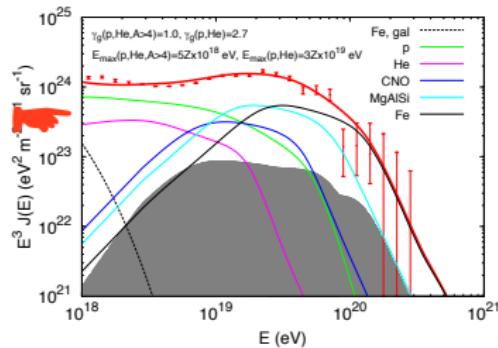
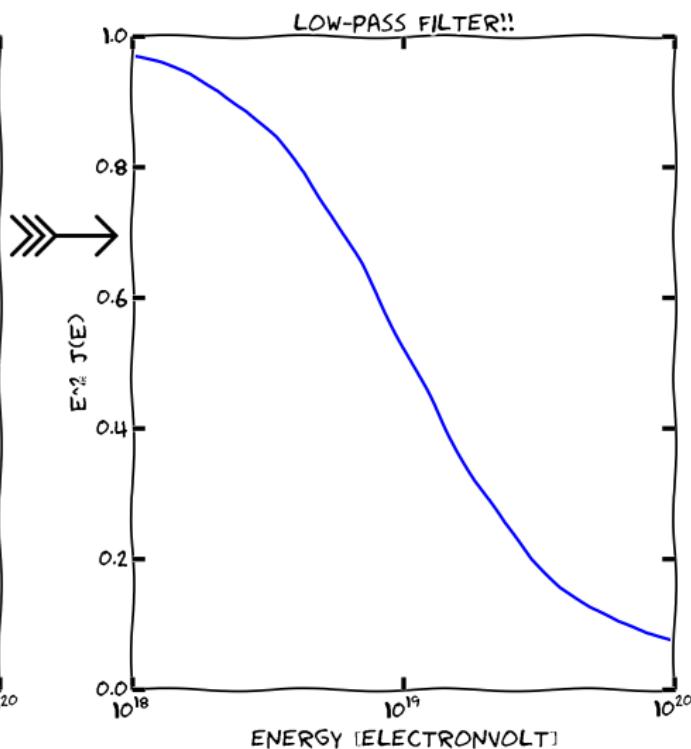
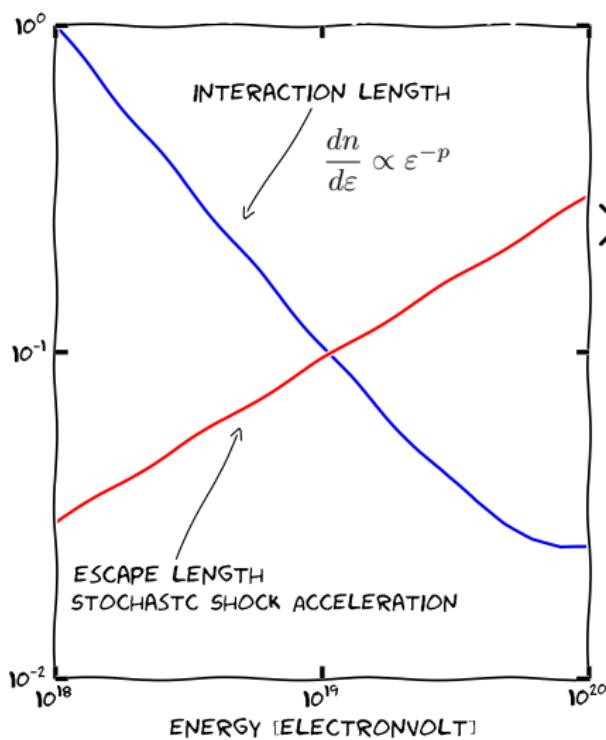
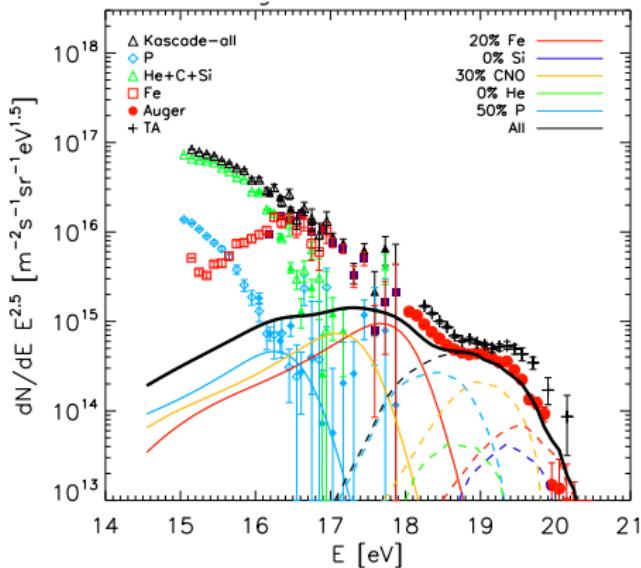
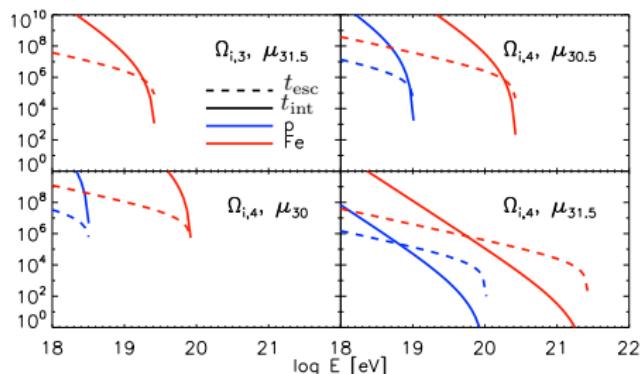


Photo-nuclear interactions during acceleration



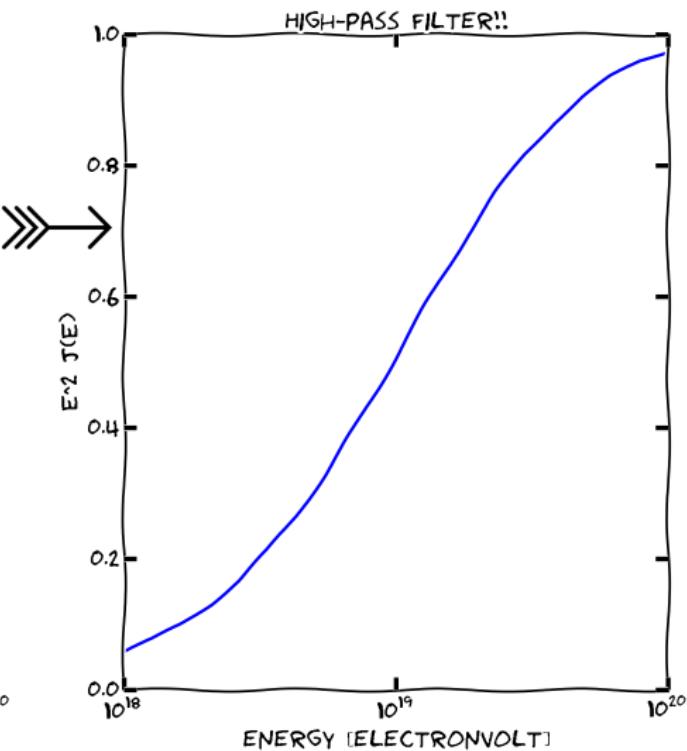
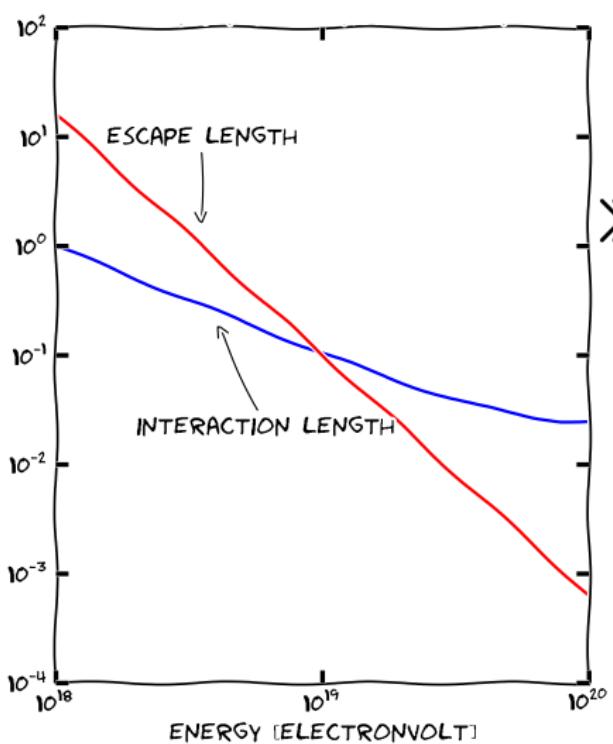
Example (Fang-Kotera-Olinto, 2012-2014)



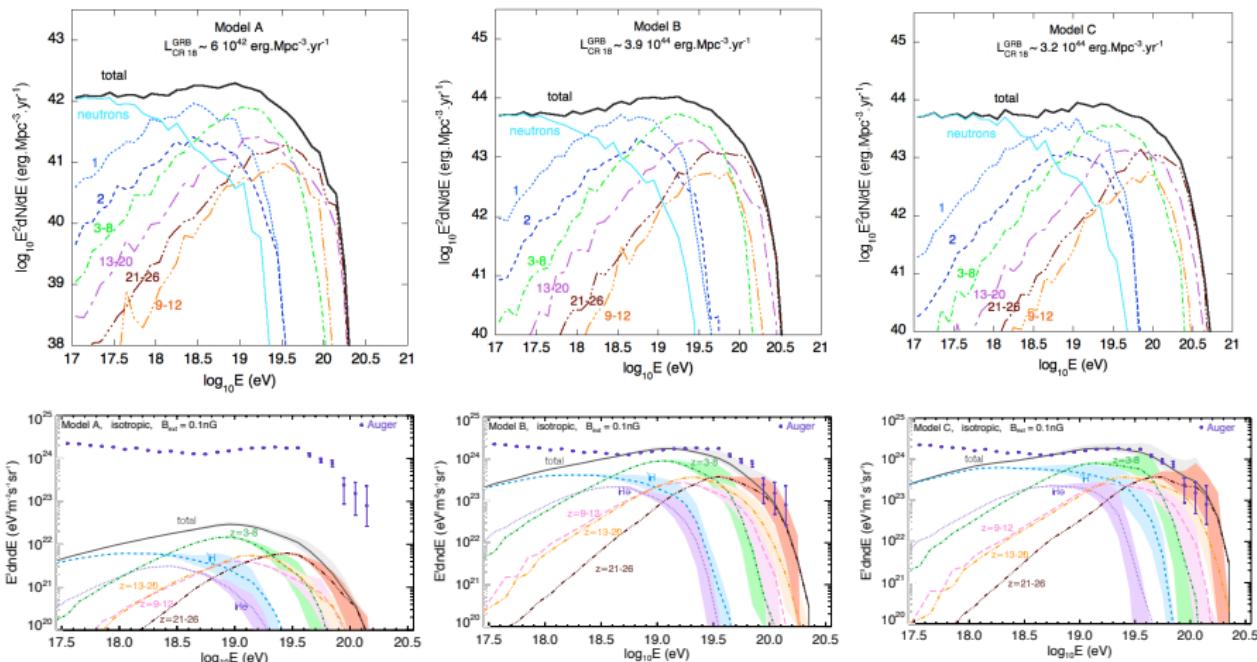
Newly-born pulsars embedded in core-collapse supernovae

- initial rotational velocity: $\Omega_i = 10^3, 10^4 \text{ s}^{-1}$
- magnetic dipole moment: $\mu = 10^{30}, 10^{31.5} \text{ cgs}$
- need GCR population (solid lines) to fill-in spectrum below ankle

New idea: photodisintegration after acceleration



Example that doesn't create ankle (Globus+, 2014)



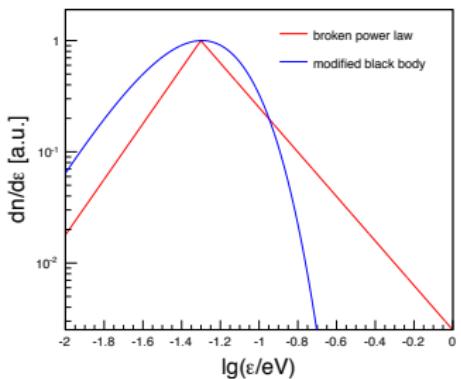
Still requires some new EG component

(Globus-Allard-Mochkovitch-Parizot, 2014)

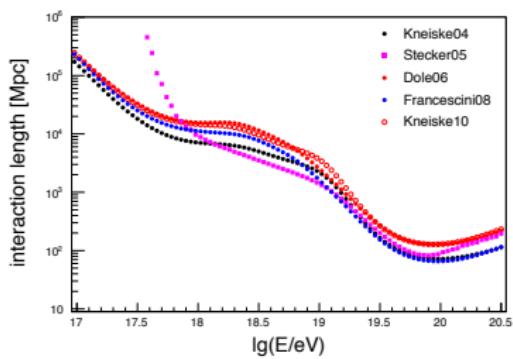
Our model: Photodisintegration in medium outside the accelerator

- injection spectrum $\propto E^{-\gamma}$
- mass of injected nuclei: A
- UHECR power density: $\dot{\epsilon}$
- source evolution with z SFR
- interaction/escape time: $t_{\text{int}} / t_{\text{esc}}$
- maximum energy: E_p^{\max}

Uncertainties

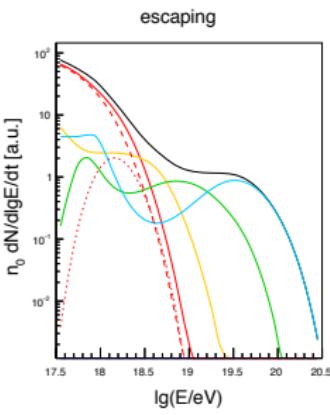
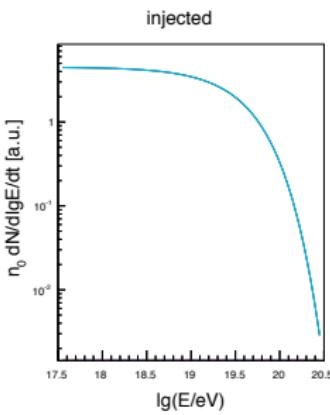
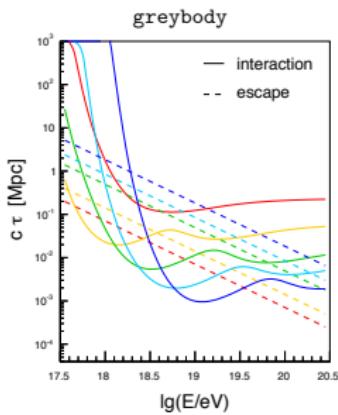
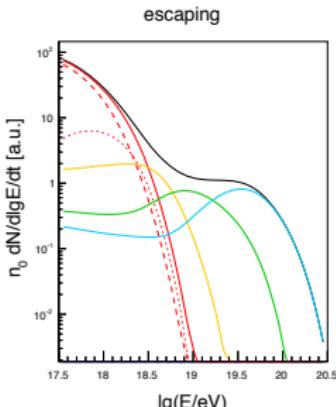
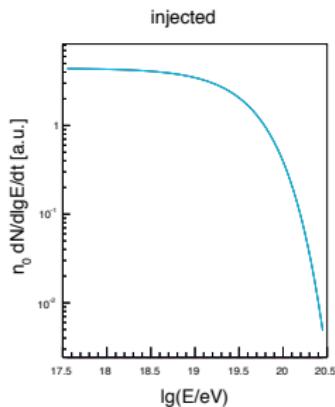
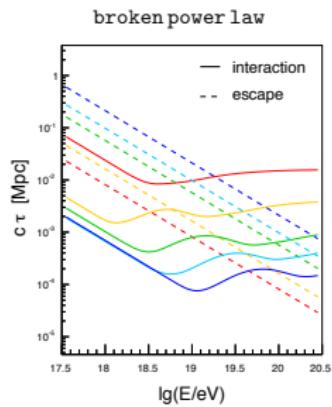


Photon field surrounding source

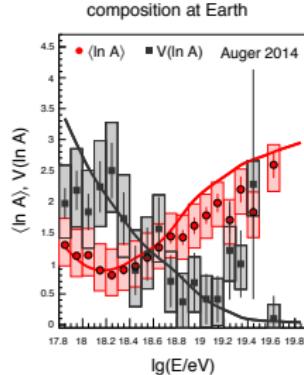
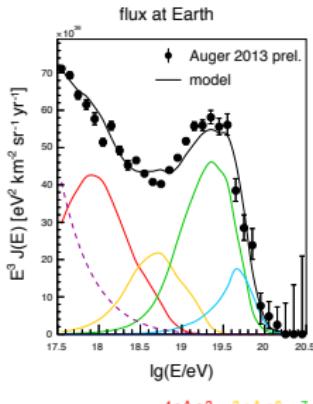


EG background light

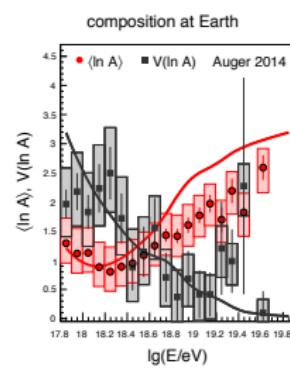
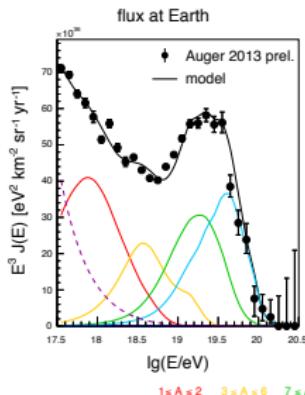
Impact of source environment depends on photon field



Example fit: broken power law



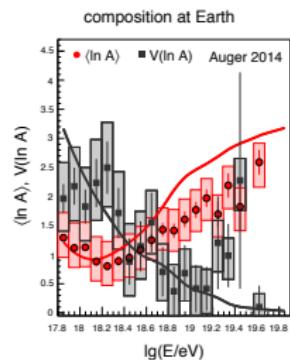
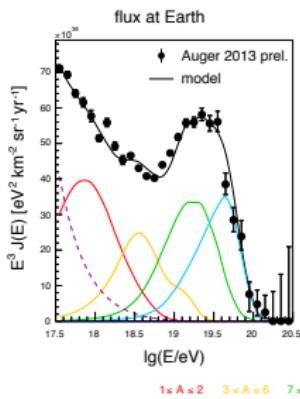
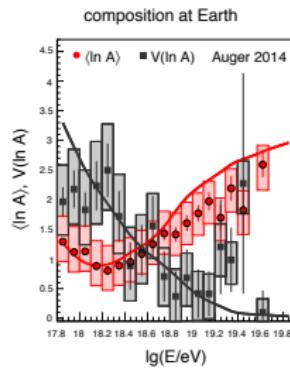
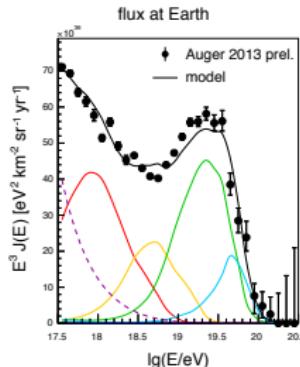
- Assuming $\gamma = 1$, Kneiske10
- $\log(E_p^{\max} / \text{eV}) = 18.6$
- $A = 23$



- Assuming $\gamma = 1$, Kneiske04
- $\log(E_p^{\max} / \text{eV}) = 18.5$
- $A = 28$

tail of GCRs (dashed line)

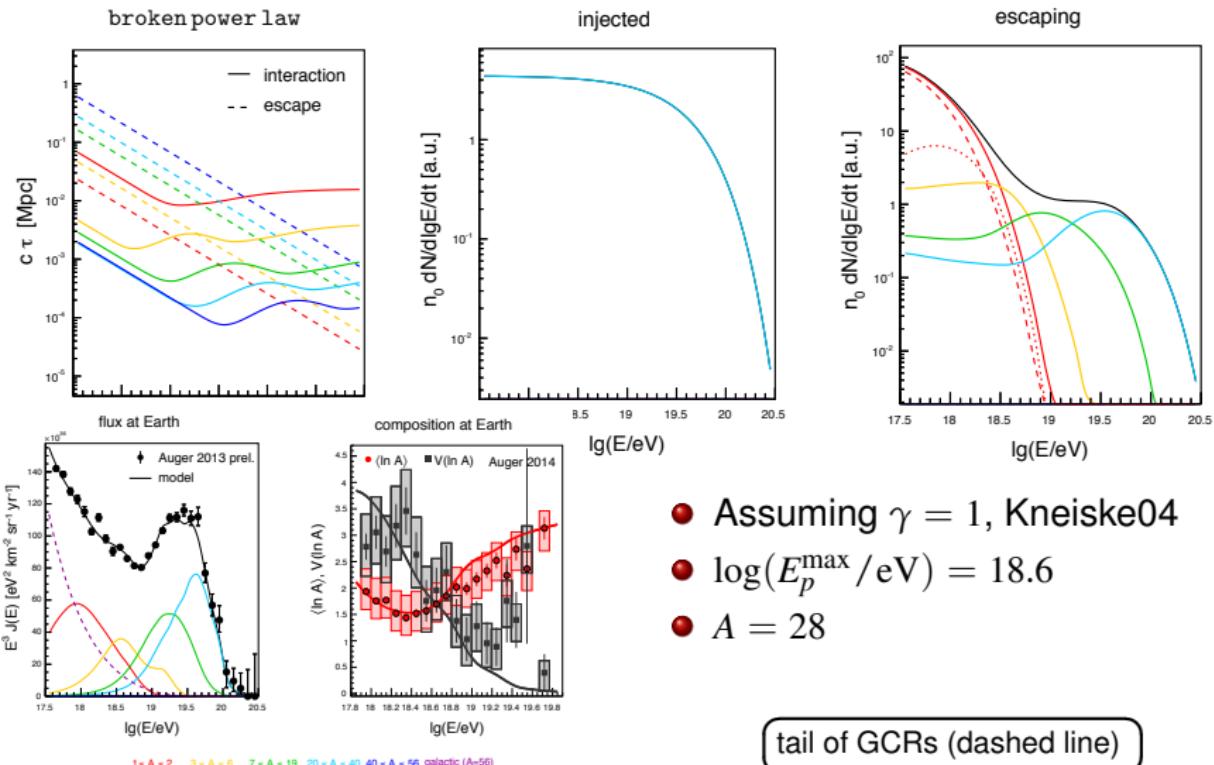
Example fit: greybody

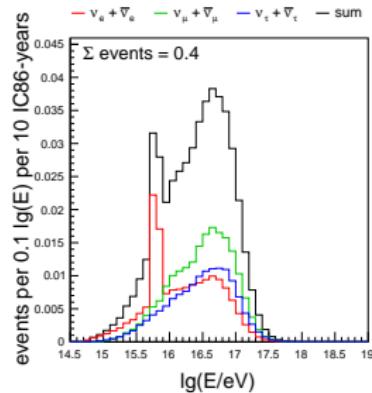
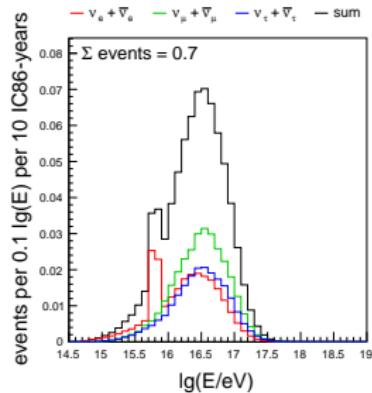
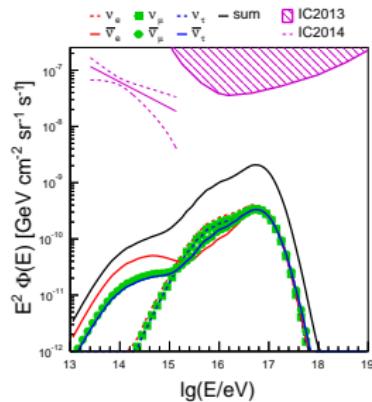
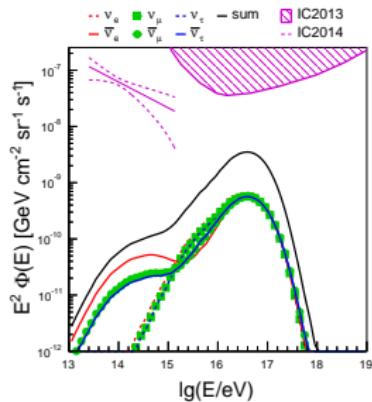


- Assuming $\gamma = 1$, Kneiske10
- $\log(E_p^{\max}/\text{eV}) = 18.6$
- $A = 23$

- Assuming $\gamma = 1$, Kneiske04
- $\log(E_p^{\max}/\text{eV}) = 18.5$
- $A = 27$

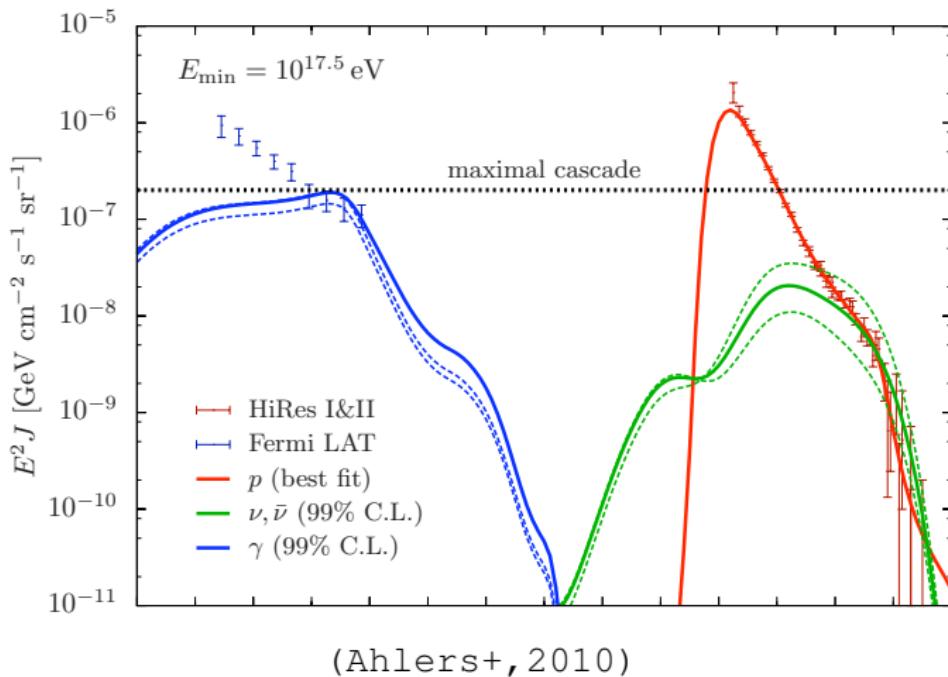
tail of GCRs (dashed line)

Systematic sensitivity (spectrum $1\sigma \uparrow$ and $\langle X_{\max} \rangle 1\sigma \downarrow$)

Distinctive ν signal broken power law (left) and greybody (right)

EM cascades from γ -rays

- Cosmogenic neutrinos safely below upper limit set by *Fermi-LAT* flux



- γ 's from nucleus de-excitation \rightleftarrows negligible contribution to EM cascades
(Anchordoqui, 2015)

Take home message

- Ankle and light extragalactic CRs below it, can be explained by photodisintegration of UHECRs *in region surrounding accelerator*
 - Auger composition *and* spectrum explained within systematics
 - Distinctive signature in neutrino flavors → testable
- Astrophysical realizations being studied – stay tuned to arXiv...